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Accepted Version

Mulholland, D. P., Lewis, S. R., Read, P. L., Madeleine, J.-B. and Forget, F. (2016) The solsticial pause on Mars: 2 modelling and investigation of causes. *Icarus*, 264. pp. 465-477. ISSN 0019-1035 doi: 10.1016/j.icarus.2015.08.038 Available at <https://centaur.reading.ac.uk/43041/>

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Published version at: <http://www.sciencedirect.com/science/article/pii/S0019103515004169>

To link to this article DOI: <http://dx.doi.org/10.1016/j.icarus.2015.08.038>

Publisher: Elsevier

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The solsticial pause on Mars: 2 Modelling and investigation of causes

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Number of pages: 43

Number of tables: 1

Number of figures: 15

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Proposed Running Head:

Modelling the martian solsticial pause

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ABSTRACT

The martian solsticial pause, presented in a companion paper (Lewis et al., this issue), was investigated further through a series of model runs using the UK version of the LMD/UK Mars Global Climate Model. It was found that
5 the pause could not be adequately reproduced if radiatively active water ice clouds were omitted from the model. When clouds were used, along with a realistic time-dependent dust opacity distribution, a substantial minimum in near-surface transient eddy activity formed around solstice in both hemispheres. The net effect of the clouds in the model is, by altering the thermal
10 structure of the atmosphere, to decrease the vertical shear of the westerly jet near the surface around solstice, and thus reduce baroclinic growth rates. A similar effect was seen under conditions of large dust loading, implying that northern midlatitude eddy activity will tend to become suppressed after a period of intense flushing storm formation around the northern cap edge.
15 Suppression of baroclinic eddy generation by the barotropic component of the flow and via diabatic eddy dissipation were also investigated as possible mechanisms leading to the formation of the solsticial pause but were found not to make major contributions. Zonal variations in topography were found to be important, as their presence results in weakened transient eddies around
20 winter solstice in both hemispheres, through modification of the near-surface flow. The zonal topographic asymmetry appears to be the primary reason for the weakness of eddy activity in the southern hemisphere relative to the northern hemisphere, and the ultimate cause of the solsticial pause in both hemispheres. The meridional topographic gradient was found to exert a much
25 weaker influence on near-surface transient eddies.

Keywords: MARS, ATMOSPHERE; MARS, CLIMATE; ATMOSPHERES,

DYNAMICS

1 Introduction

In Lewis et al. (this issue), hereafter referred to as P1, martian transient waves
30 were diagnosed from a reanalysis of Thermal Emission Spectrometer (TES)
data. A particular feature identified in the record was the ‘solsticial pause’:
a weakening of waves near the surface around winter solstice of each year
analysed, seen in both hemispheres.

Solsticial pauses with a range of depths have been simulated by martian global
35 climate models (MGCMs) (Hourdin et al., 1995; Basu et al., 2006; Wang et al.,
2013; Kavulich Jr et al., 2013), and existing modelling literature suggests that
the depth of a modelled solsticial pause is enhanced by the presence of one or
both of a large solsticial dust loading (representative of a global dust storm)
(Hourdin et al., 1995; Kuroda et al., 2007) and radiatively active water ice
40 clouds (Wilson, 2011). The reduction in eddy activity under conditions of high
dust loading is understood to be a result of changes to the zonal jet (Kuroda
et al., 2007), and several important features of winter hemisphere eddies have
been drawn from observational data (Wang, 2007), but a detailed description
of the mechanisms behind the solsticial pause, which occurs during years both
45 with and without a global dust storm, is currently lacking.

The midwinter minimum in North Pacific atmospheric storminess on Earth,
though different in several ways to the martian solsticial pause (see P1), has
been studied in detail in recent years, and some of the explanations put forward
for the terrestrial case may be relevant to Mars. These include: an increase in
50 barotropic damping around solstice (Deng and Mak, 2006); eddy dissipation
through diabatic effects (Chang, 2001); and localised effects of interaction with

topography (Penny et al., 2010; Park et al., 2010).

In this paper we use an MGCM to measure the net effects of several of these factors, some of which may contribute to the formation of the pause, in both
55 hemispheres. First, in section 2, the model used is described. In section 3, the set of simulations are introduced, along with summary measures of the extent to which each is able to reproduce solsticial minima similar to those shown in P1. Section 4 analyses the contributions of various mechanisms to the formation of the solsticial pause, from seasonal variation in atmospheric
60 baroclinicity, to other means of reducing eddy growth, and finally to the more fundamental impact of surface topography. In section 5, the role of topography is discussed further, and our insights into the nature of the pause are used to provide explanation for its interannual variability as presented in P1. Finally, the key results of the paper are summarised in section 6.

65 **2 Model description**

The model used from this study is an updated, UK version of the LMD/UK Mars Global Climate Model (UKMGCM) described in P1. This version includes a microphysical cloud scheme, which predicts ice particle sizes and growth rates taking into account temperature, humidity and local density of
70 dust nuclei (Montmessin et al., 2004; Madeleine et al., 2012). Water vapour and ice tracers undergo parameterised turbulent diffusion, convective adjustment and gravitational sedimentation. The radiative effect of water ice clouds can be included; the radiation scheme (Madeleine et al., 2011) calculates cloud extinction depending on the ice particle size, using optical parameters from
75 Warren (1984) (see Madeleine et al. (2012) for more details).

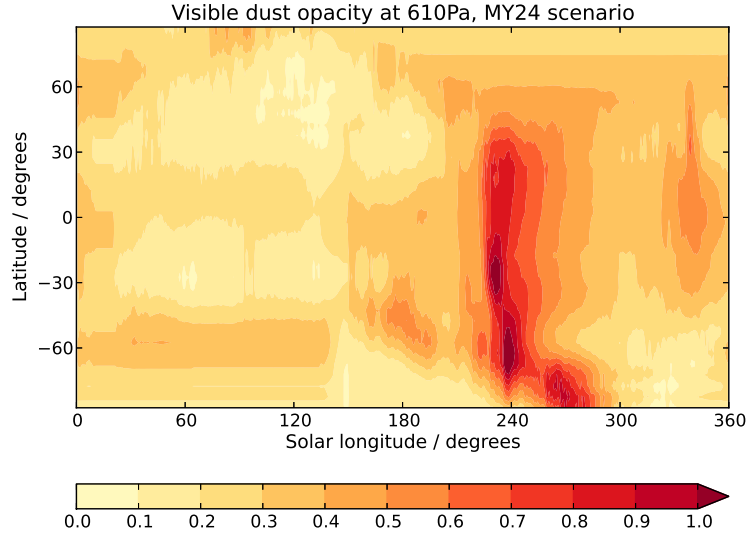


Fig. 1. The ‘MY24’ dust scenario, constructed from a combination of MY24 and MY25 reanalysis dust opacities, plotted as a zonal mean function of latitude and solar longitude.

Full lifting and transport simulations using radiatively active dust are possible in the UK model, but in this work, prescribed dust opacities are used in lieu of a transported dust field, to enable more direct comparisons with observations and to fix the dust opacity field within a set of controlled model experiments
80 used to study the sensitivity of the solsticial pause to several different effects. In this mode, the dust field that the radiation scheme ‘sees’ is comprised of particles of radius $1.5\mu\text{m}$, with variation of mass mixing ratio with height following a Conrath profile. Mixing ratios are scaled so that the total column optical depth, calculated in two visible/near-IR bands and three thermal IR
85 bands and using a ratio of visible to infrared opacity of 2, fits a prescribed distribution. The visible optical depth may have a constant value in time and space, or follow a more complex spatially and/or temporally varying function (Forget et al., 1999).

One such distribution used in this work varies with longitude, latitude and
90 time, and replicates dust opacities from MY24 as observed by the Thermal
Emission Spectrometer onboard Mars Global Surveyor (Smith et al., 2000). It
was constructed using the reanalysis dataset described in P1. Since data were
not available for $L_s = 0\text{--}140^\circ$ in MY24, the dust opacity scenario for this year
uses optical depths from MY25 of the reanalysis dataset during this period.
95 The differences in dust loading between years is expected to be fairly small
at this time of the martian year (northern spring/summer). The ‘MY24’ dust
scenario is shown as a zonal mean in Fig. 1.

The key point is that model runs using the ‘MY24’ scenario can be compared
directly to the reanalysis. Aside from possible inaccuracies in the vertical dust
100 distribution, which will be compensated for in the reanalysis (through temper-
ature assimilation) but not in the free-running model, any differences between
model and reanalysis represent deficiencies in the model’s ability to simulate
correctly the atmospheric state (temperature, pressure, wind) given the cor-
rect dust opacity forcing.

105 Note that there is considerable uncertainty in the column dust opacity in both
winter polar regions, as there were no TES opacity data available in these
regions. As a result of the approach taken in the data assimilation procedure,
dust opacities remain constant over a period without available observations,
as can be discerned from Fig. 1. The simulations presented in the next section
110 include the use of two alternative prescribed dust distributions, which gives
information on the sensitivity to dust loading more generally, and provides
some confidence that the impact of water ice clouds is not specific to the use
of a particular, possibly inaccurate, polar dust distribution.

The model runs described below were all carried out using the T31 spectral
truncation, corresponding to a $3.75^\circ \times 3.75^\circ$ dynamical grid and a $5^\circ \times 5^\circ$
physics grid, the same as was used for the reanalysis of P1. The use of a free-
running MGCM complements the analysis of observational and assimilated
data by allowing the isolation of the net effects of each of several physical
mechanisms, through the use or neglect of each in a particular simulation.
In this work, the three model forcings that were varied in this way were the
dust opacity field, the radiative effects of water ice clouds and the surface
topography.

3 Model simulations

To test whether or not the model reproduces the reanalysis as presented in
P1, and to investigate the conditions necessary for the formation of a solsticial
pause, a series of simulations were performed using, in each case, a prescribed
dust opacity field. The simulations are named and summarised in Table 1. Two
used the ‘MY24’ dust scenario described previously, one including the radiative
effects of water ice clouds (τ_{MY24}^*) and the other neglecting clouds (τ_{MY24}).
A similar pair of runs with (τ_{low}^*) and without (τ_{low}) clouds used instead a
constant visible dust opacity of 0.2, referenced to the 610 Pa pressure surface.
Finally, τ_{high} simulated a large, perennial dust storm by using a constant
610 Pa dust opacity of 2.0, and neglected water ice clouds.

Table 1 also includes a quantity $\frac{[T'_{max}]_{solstice}}{[T'_{max}]_{pre.post}}$ (along with its constituent parts)
which measures, in a crude way, the extent to which a solsticial pause was
produced in the northern hemisphere for a particular run. This quantity was
calculated by finding the maximum value of the RMS transient temperature,

filtered to include only signals with period between 1.5 and 30 sols, at a height of 2.5 km within the region 30–90°N, at each point in time, and taking the ratio of its averages over $L_s = 240\text{--}300^\circ$ and over $L_s = (180\text{--}240^\circ, 300\text{--}360^\circ)$. Low values therefore signify a deep minimum in transient eddy activity: the reanalysis value for MY24 is 0.50.

In summary, all runs except τ_{low} produced a discernible solstitial minimum in near-surface eddy activity. τ_{MY24}^* came closest to replicating the results of the reanalysis, though fell somewhat short. A deeper solstitial pause was produced whenever dust optical depth was increased (compare τ_{low} , τ_{MY24} and τ_{high}), or whenever clouds were activated. The results suggest that both the presence of water ice clouds and an increase in global dust loading close to solstice contribute to the development of a solstitial pause.

Name	Dust opacity	Clouds?	$[T'_{max}]_{solstice}$	$[T'_{max}]_{pre,post}$	$\frac{[T'_{max}]_{solstice}}{[T'_{max}]_{pre,post}}$
τ_{MY24}	MY24	No	6.40	7.09	0.90
τ_{MY24}^*	MY24	Yes	5.55	7.94	0.70
τ_{low}	0.2	No	7.91	7.42	1.07
τ_{low}^*	0.2	Yes	8.39	9.18	0.91
τ_{high}	2.0	No	5.51	7.27	0.76
MY24 reanalysis			3.35	6.72	0.50

Table 1

Summary of the model simulations performed.

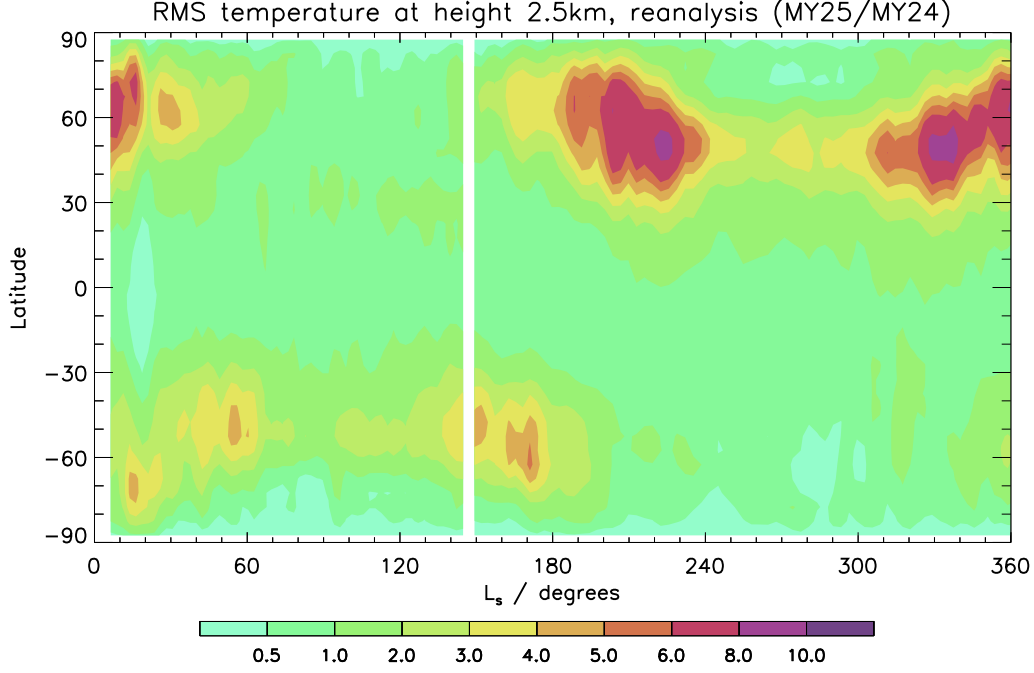


Fig. 2. RMS transient temperature (K) at a height of 2.5 km above the surface, from the reanalysis. Shown is a composite of MY25 (for $L_s = 0\text{--}140^\circ$) and MY24 ($L_s = 140\text{--}360^\circ$), which is the same split as in the ‘MY24’ dust scenario used below (in model simulations τ_{MY24} and τ_{MY24}^*). Values are masked out for one sol around the split at $L_s = 140^\circ$.

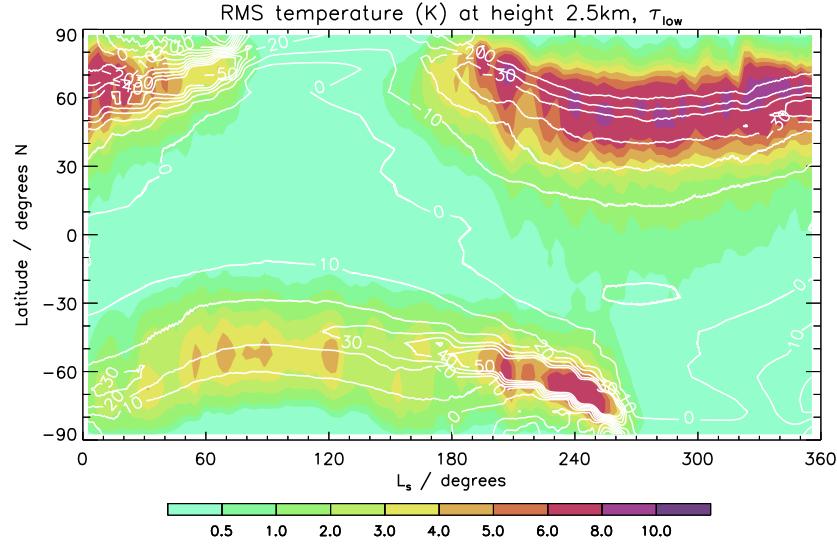
150 To aid comparison with model simulations, one year of zonal-mean RMS tran-
 sient temperature at 2.5 km above the surface from the reanalysis (specifically
 a hybrid of MY24 and MY25, formed in the same way as the ‘MY24’ dust
 opacity scenario, described above) is displayed in Fig. 2. The same quantity for
 the full reanalysis record can be seen in P1, Fig. 1. In this and all subsequent
 155 plots, temporal filtering is used to retain only eddies with period between 1.5
 and 30 sols, which are termed ‘transient’, and time series are smoothed using
 an averaging window of length 9 sols. The solsticial pause is clearly visible in
 each hemisphere, as is the stronger peak amplitude of transient eddies in the
 northern hemisphere than in the southern hemisphere.

160 3.1 Without water ice clouds (τ_{low} , τ_{MY24})

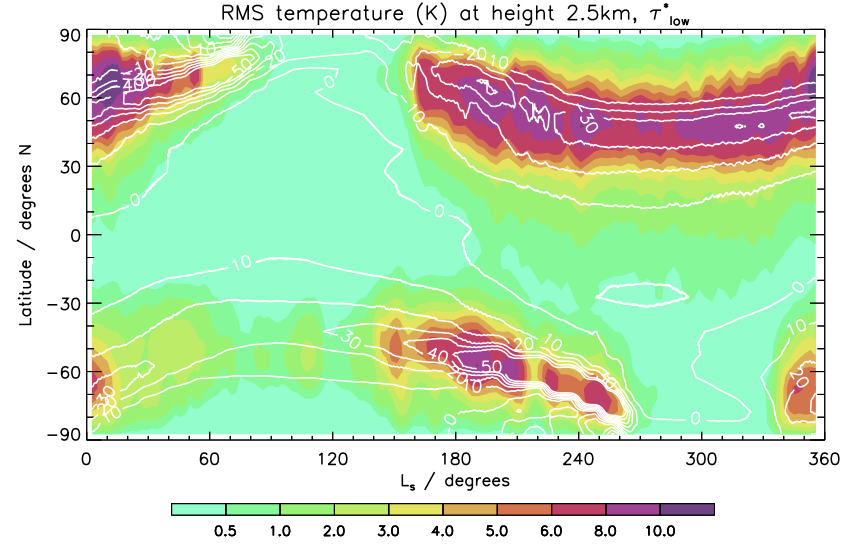
Fig. 3 (left column) shows the equivalent RMS transient temperature from model simulations τ_{low} and τ_{MY24} , both performed without the radiative impact of water ice clouds. The constant opacity case shows no sign of a solsticial pause in the northern hemisphere, with a peak RMS of 8–9 K throughout the period $L_s = 180$ – 360° . With the MY24 dust scenario, a minimum in RMS develops between $L_s = 260^\circ$ and $L_s = 290^\circ$, but it is much smaller in amplitude and shorter in duration than that seen in the reanalysis. The overlaid meridional temperature gradients show that in both cases the peak in transient eddy activity lies around 5° to the north of the latitude where the temperature gradient is strongest. Evidently the MY24 dust scenario, with elevated opacities over most of the planet from $L_s \sim 210^\circ$ onwards, goes some way towards forcing a solsticial pause, perhaps for a similar reason to that identified by Kuroda et al. (2007) for a large solsticial dust storm; however, this realistic dust loading falls well short of reproducing the results from the reanalysis for the same year. The southern hemisphere shows similar results, with fairly constant eddy activity for τ_{low} and a short (though otherwise quite realistic) minimum for τ_{MY24} .

3.2 With active water ice clouds (τ_{low}^* , τ_{MY24}^*)

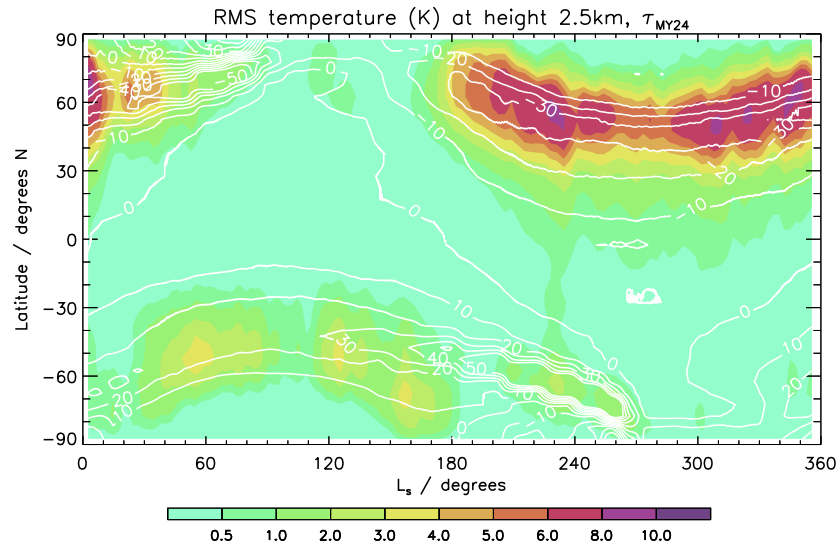
Fig. 3 (right column) shows the results of a similar pair of runs but with radiatively active clouds now included (τ_{low}^* and τ_{MY24}^*). Focusing again on the northern hemisphere, with MY24 dust opacity the solsticial pause is now wider and deeper than before, and resembles somewhat the observed pause



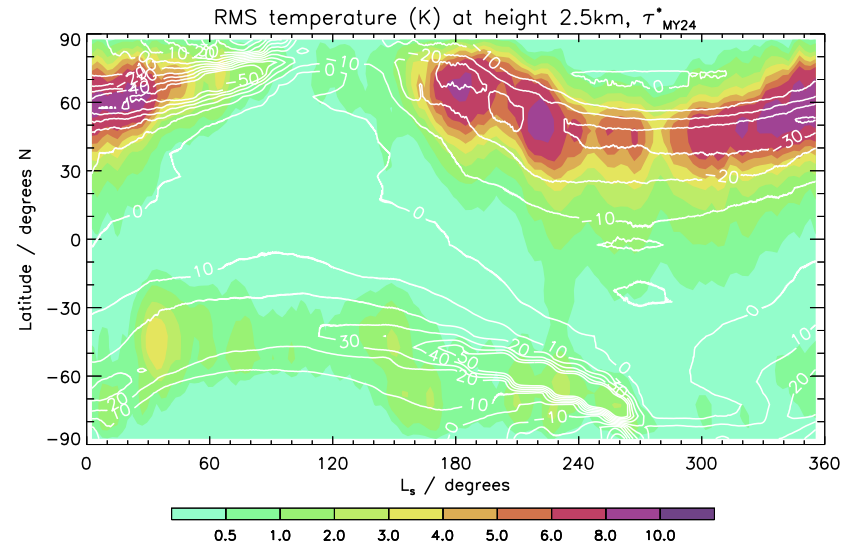
(a)



(b)



(c)



(d)

Fig. 3. RMS transient temperature (K) at a height of 2.5 km above the surface, as in Fig. 2, output from the MGCM from simulations (a) τ_{low} , (b) τ_{low}^* , (c) τ_{MY24} and (d) τ_{MY24}^* (moving from top to bottom, the more realistic dust distribution is introduced, and moving from left to right, radiatively active ice clouds are introduced). Meridional temperature gradient $\frac{dT}{dy}$ at 200 m above the ground (10^{-6} K m^{-1}) is overlaid in white contours.

in the reanalysis, though the model pause is much less well defined (RMS temperature does not drop off sharply enough before $L_s = 240^\circ$), and transient eddy amplitudes are slightly larger throughout the northern autumn/winter period. Compared to τ_{MY24} , the pause has been deepened by both increased eddy activity in early autumn ($L_s = 180\text{--}220^\circ$) and late winter ($320\text{--}360^\circ$), and weakened activity around solstice. The southern hemisphere pause is also enhanced, and is now wider than that of the reanalysis. τ_{low}^* also simulates a solstitial pause in the southern hemisphere, and a weak pause in the northern hemisphere. Particularly in the pre-solstice period in τ_{low}^* , and throughout $L_s \approx 180\text{--}300^\circ$ in τ_{MY24}^* , the peak in eddy activity is now located at or to the south of the latitude of strongest temperature gradient. In τ_{MY24}^* , around solstice, the location of the peak temperature gradient is roughly unchanged from τ_{MY24} , meaning that the peak in eddy activity is shifted southward by $\sim 5^\circ$, compared to τ_{MY24} .

For τ_{MY24}^* , zonal wavenumbers 1–3 are all reduced to similar levels around solstice. Transitions between the various modes occur during the autumn/winter season. $s = 1$ perturbations tend to be confined north of 45°N , while $s = 3$ are located south of 60°N , similar to what was seen in MY26 by Wang (2007).

The cloud ice mixing ratios produced in simulation τ_{MY24}^* are shown in Fig. 4. Polar hood clouds are confined below 100 Pa, and are generally thickest near the seasonal cap edge (which, for reference, is located at approximately $(60^\circ, 50^\circ, 50^\circ)\text{N}$ at $L_s = (220^\circ, 270^\circ, 320^\circ)$). Cloud optical depth in this model run reaches a maximum around the time of a regional dust storm at $L_s = 240^\circ$, and declines thereafter. It is also notable that cloud optical depth briefly drops to a minimum at $L_s \approx 200^\circ$ (not shown), coincident with the gap between the two pre-solstice peaks seen in Fig. 3(d).

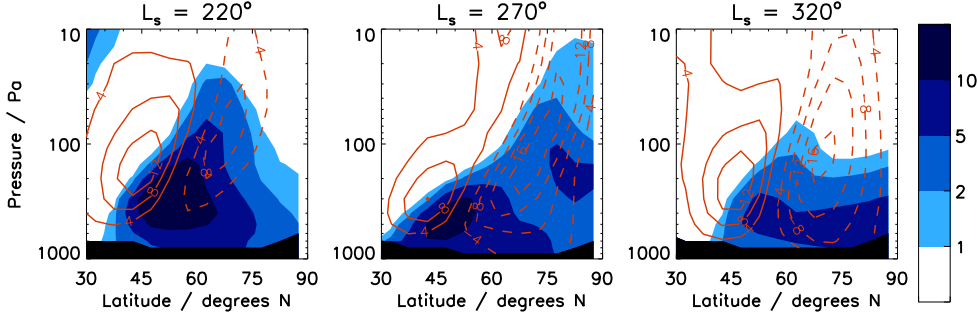


Fig. 4. Change in \bar{u} , $(\tau_{MY24}^* - \tau_{MY24})$ (m s^{-1} , contours, positive values solid, negative values dashed), and water ice cloud mixing ratio for τ_{MY24}^* (10^{-6} kg/kg , shading). The surface is coloured in black.

Compared to observations, column ice opacities in northern winter mid- and
 210 high latitudes are somewhat overestimated by the model: visible ice opaci-
 ties in the model north polar hood are typically 0.2–0.5, and reach values of
 around 1 at their peak at $L_s \approx 240^\circ$, whereas observations from Mars Climate
 Sounder (McCleese et al., 2007; Benson et al., 2011) found that typical north
 polar hood visible opacities are around 0.1–0.3. Therefore, the thermodynami-
 215 cal impact of the clouds (seen in the change to the zonal wind, drawn in Fig. 4,
 and discussed further in the next section) may be greater in the UKMGCM
 than it is in reality, and so quantitative aspects of some of the results shown
 in the next section should be considered as specific to the UKMGCM. The
 qualitative robustness of the impact of water ice clouds on the solsticial pause
 220 is discussed in section 5.

4 Mechanisms behind the pause

4.1 Changes in atmospheric baroclinicity

Having found that the inclusion or removal of water ice clouds strongly affects the ability of the MGCM to simulate the solsticial pause, the model runs were
 225 analysed more closely to determine the manner in which the clouds exert their influence. The most obvious mechanism for this is that the clouds alter the structure of the atmosphere so as to reduce baroclinicity at low levels around solstice.

4.1.1 Basic atmospheric state

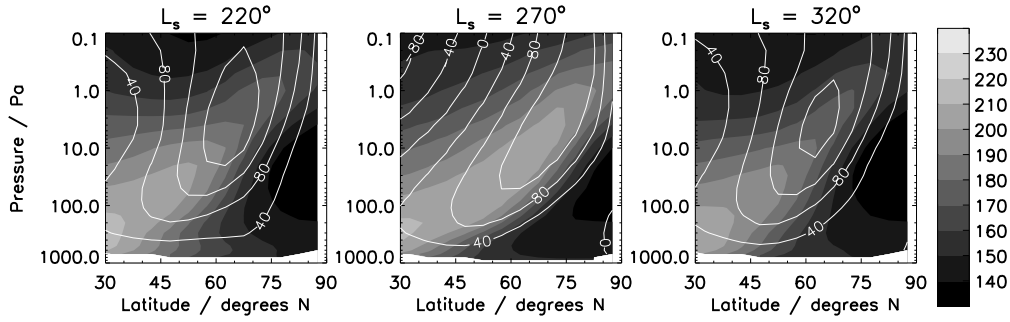


Fig. 5. Zonal-mean temperature \bar{T} (K, shading) and zonal wind \bar{u} (m s^{-1} , contours) for τ_{MY24}^* .

230 Fig. 5 shows the zonal-mean state of τ_{MY24}^* , the most realistic of the model runs. The change to the zonal-mean zonal wind below 10 Pa caused by the addition of water ice clouds can be seen in Fig. 4. The addition of clouds slows the zonal wind by $5\text{--}10 \text{ m s}^{-1}$ near the ground at $60\text{--}80^\circ\text{N}$ around solstice, and below 100 Pa the jet is shifted equatorward in such a way as to enhance
 235 the poleward tilt of this lower portion. The change to the thermal structure

near the ground around solstice encourages this, as temperature contours at 40–70°N (Fig. 5) are roughly aligned with the vertical before and after solstice but slope strongly around solstice as the surface at 40–70°N cools. Since cloud formation is determined primarily by atmospheric temperature (Benson et al., 2011), the clouds adopt approximately the same sloped structure around solstice (Fig. 4). A cooling of up to 5–10 K occurs in the vicinity of and below the clouds, the effect of which is to reduce the meridional temperature gradient on the clouds’ poleward side, and strengthen it on the equatorward side. The clouds do, therefore, alter the mean state of the atmosphere, and so have the potential to affect baroclinic instability.

In the MY24 reanalysis, the reduction in zonal wind near the surface at 60–80°N begins earlier in the year than it does in τ_{MY24}^* , along with the more rapid reduction in eddy activity seen previously (compare Figs. 2 and 3(d) at $L_s \approx 230^\circ$).

4.1.2 Near-surface temperature gradients

A comparison of the left and right columns of Fig. 3 shows that the maximum (most negative) meridional temperature gradient very close to the ground in the northern hemisphere is weakened by the addition of clouds (seen particularly during the pre-solstice period), since the peak in cloud opacity tends to lie to the south of the peak temperature gradient, implying a reduction to the peak temperature gradient (on the clouds’ poleward side) as explained above. The latitude of the peak gradient is also shifted $\sim 5^\circ$ to the north before solstice in τ_{MY24}^* relative to τ_{MY24} , but reaches a similar latitude of ~ 45 – 50° N at solstice in both τ_{MY24} and τ_{MY24}^* , in good agreement with the reanalysis.

260 Transient eddy activity does not closely follow the strength of the temperature gradient; for example, compare the very similar gradients seen in southern midlatitudes at $L_s = 120\text{--}210^\circ$ in all four simulations shown in Fig. 3, with the different temperature variances that result. The reanalysis itself (not shown) features a northern midlatitudes temperature gradient at height 200 m that is
 265 fairly constant over $L_s = 220\text{--}310^\circ$, and is stronger at solstice than at the time of the pre-solstice maximum ($L_s = 200\text{--}220^\circ$), though it does increase around the time of the post-solstice maximum. Therefore, near-surface baroclinic eddy activity cannot simply be explained in terms of the varying magnitude of the near-surface temperature gradients.

270 4.1.3 Linear growth rates

A better measure of the potential for baroclinic instability of the lower atmosphere is provided by the Eady growth rate, σ (e.g. James and Gray (1986)):

$$\sigma = 0.31 \frac{f}{N} \frac{\partial \bar{u}}{\partial z}$$

where f is the Coriolis parameter, N is the atmospheric static stability and \bar{u} is the zonal-mean zonal wind velocity. Since the baroclinic zone moves equatorward from autumn to winter with the growth of the seasonal CO_2 cap, the change in f ($\propto \sin(\text{latitude})$) will cause σ to decrease by a factor of ~ 1.3
 275 from 70° to 45° latitude. In addition, Wang et al. (2005) found that in the TES data, N^2 , at a height of 2.5 km above the north polar cap edge, peaked at $L_s = 260^\circ$, at around twice its value at $L_s = 200^\circ$, implying an increase in N by a factor of ~ 1.4 from autumn to winter (and these results were fairly well reproduced by the model, both with and without ice clouds). Therefore,
 280 simply following the edge of the seasonal polar cap, basic state conditions

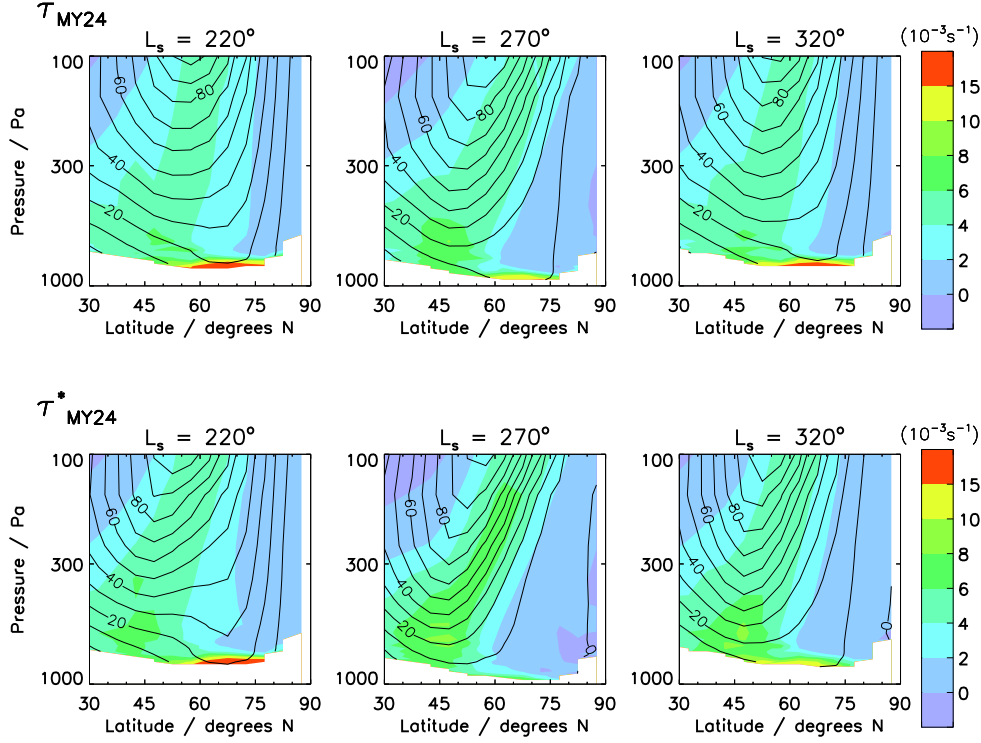


Fig. 6. Zonal mean vertical wind shear $\frac{\partial \bar{u}}{\partial z}$ (shading, s^{-1}) and zonal wind \bar{u} , (contours, ms^{-1}) near the surface, for runs τ_{MY24} (top) and τ_{MY24}^* (bottom).

suggest a reduction in baroclinic growth rate of around 40% from autumn to winter solstice. However, the vertical wind shear, $\frac{\partial \bar{u}}{\partial z}$, has a crucial role in determining the growth rates that ultimately occur before and at solstice.

The growth rate in τ_{MY24}^* exhibits a maximum around solstice at 20 Pa,
 285 $\sim 70^\circ\text{N}$, in agreement with the simulated middle atmosphere RMS temperature peak (not shown). The addition of clouds, most importantly at high altitudes in the tropics, increases this upper maximum in comparison with τ_{MY24} , by strengthening the principal meridional overturning cell (PMOC), and increasing the polar warming and the associated vertical wind shear. Growth
 290 rate σ also maximises at solstice at $\sim 50^\circ\text{N}$ at altitudes as low as 400 Pa.

However, the growth rates that are important with respect to the solsticial pause are situated very close to the ground, within the lowest few hundred metres, and vary partly due to changes in vertical wind shear. Without clouds (τ_{MY24}), a relatively large surface wind shear of $\geq 0.01 \text{ s}^{-1}$ exists at 60–80°N throughout the period shown (Fig. 6 (top)). It is slightly reduced around solstice, compared to its values in late autumn and early spring. When clouds are used, this reduction in shear around solstice is more substantial, to $\approx 0.006 \text{ s}^{-1}$ (Fig. 6 (bottom)). The jet also appears sharper and tilted further from the vertical around 100 Pa, with an equatorward shift closer to the ground (note in particular the shift in the 10 m s^{-1} contour), at $L_s = 270^\circ$.

This correlation between Eady growth rate and RMS transient eddy temperature (which was also noted by Wang et al. (2013), using a slightly different baroclinic parameter) is clearly displayed in Fig. 7, which shows the maximum values of the two quantities (growth rate and RMS temperature at 200 m and 2.5 km above the surface, respectively) within 30–90°N, along with the growth rate (at 200 m height) at the location of the eddy temperature maximum. For τ_{MY24}^* (middle panel), the peak RMS temperature (solid line) in the northern domain shows a correlation with the maximum Eady growth rate (dashed line), and both fall to half their $L_s = 210^\circ$ values by $L_s = 280^\circ$. However, the locations of the two maxima differ somewhat after $L_s \approx 210^\circ$. As can be inferred from Fig. 6 (bottom panel), the peak Eady growth rate is found at 60–80°N throughout the season, whereas the maximum in RMS temperature is centred south of 60°N after about $L_s = 210^\circ$. The post-solstice increase in eddy activity ($L_s = 290\text{--}350^\circ$) is not matched by a significant local rise in growth rate (dotted line), but rather by an increase in the growth rate to the north of this location, at $\sim 70^\circ\text{N}$.

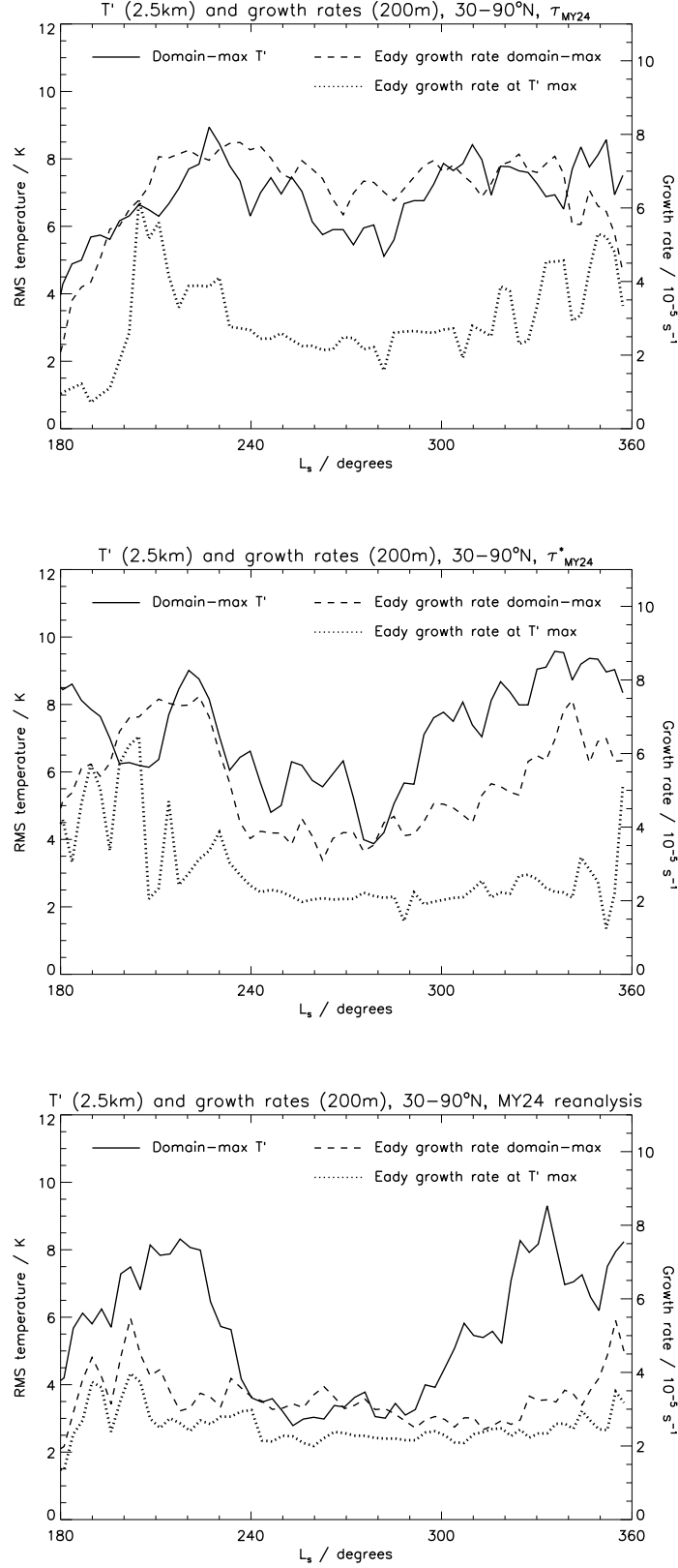


Fig. 7. Peak values of the zonal-mean RMS transient eddy temperature at 2.5 km height (solid, left axis) and Eady growth rate at 200 m height (dashed, right axis) within 30–90°N, and zonal-mean 200 m Eady growth rate at the latitude of the peak in RMS temperature (dotted, right axis), during northern autumn and winter, for τ_{MY24} (top), τ_{MY24}^* (middle) and the reanalysis of MY24 (bottom).

The same plot for τ_{MY24} (top panel) shows that its more moderate solstitial RMS temperature reduction is accompanied by reductions in both local and domain-maximum Eady growth rates. The local growth rate at solstice again
320 falls to less than half of its pre-solstice value, suggesting that the maximum RMS temperature, which only decreases to around two thirds of its pre-solstice value, is to some extent being maintained by the presence of a larger growth rate, situated to the north, which displays only a slight solstitial minimum. This could be possible through equatorward propagation of disturbances from
325 60–70°N, since the meridional flow near the surface is relatively strong (along the topographic channels) in the autumn period, as evidenced by the propagation of flushing dust storms in the same direction.

In the reanalysis (bottom panel), the situation appears more complicated, and RMS temperature peaks are quantitatively less well predicted by the
330 Eady growth rate. While the maximum RMS temperature falls to less than half of its peak value at solstice, the local growth rate decreases by a smaller fraction. The latitude of peak growth rate coincides more closely with peak eddy activity than it does in the model (seen in the similarity of the dashed and dotted lines), as peak growth rates move southward to $\sim 60^\circ\text{N}$ over $L_s = 220$ –
335 330° . The post-solstice rise in RMS temperature after $L_s = 320^\circ$ is matched by a modest rise in growth rate, as the CO_2 cap begins to recede and the latitude of peak eddy activity moves slightly northward. It should be noted that the quasi-geostrophic assumption behind the Eady growth rate is likely to be tested at such low altitudes, within the turbulent planetary boundary
340 layer. Nevertheless, these results may suggest that the real situation is more complex than the dynamics being simulated by the MGCM, even when ice clouds are included.

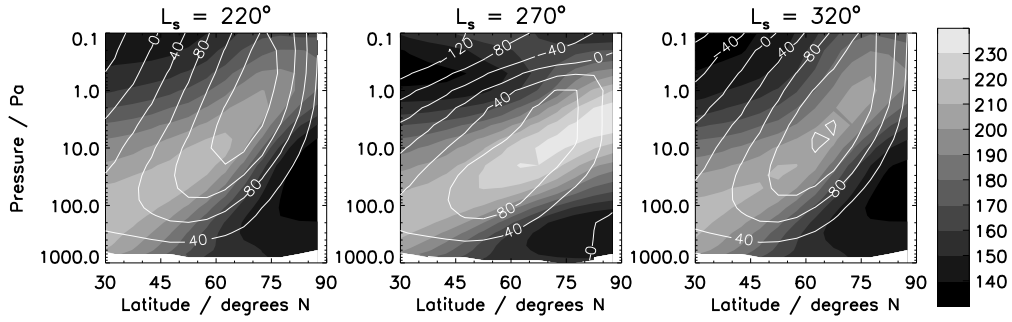


Fig. 8. As the top panel of Fig. 5, but for τ_{high} .

4.1.4 Similarities with ‘dust storm’ simulation (τ_{high})

Similar results were observed in the dust storm simulation τ_{high} , which produced a significant northern hemisphere pause (see Table 1) despite neglecting water ice clouds. As shown in other studies, an increased dust loading enhances the mean meridional circulation and strongly increases the polar warming over northern high latitudes (as much as 60 K warming at the 10 Pa level, in comparison to τ_{low}) (Fig. 8). This results in a stronger tilt of the entire jet, which leads to a reduction in vertical shear near the surface at 60°N, similar to that in τ_{MY24}^* . It is at winter solstice that the PMOC is strongest and the tilt of the jet most pronounced, and thus the instability at the ground is strongly reduced. $s = 1$ waves are most strongly affected, and $s \geq 3$ less so, giving a decrease in the dominant wavelength, in agreement with Barnes et al. (1993) and Kuroda et al. (2007).

In this simulation the variation in peak RMS temperature is well matched by the local Eady growth rate, as both drop to around one third of their pre-solstice values at solstice (Fig. 9). The maximum growth rate within the domain also reaches a minimum, at $L_s = 280^\circ$, of around one third of its pre-

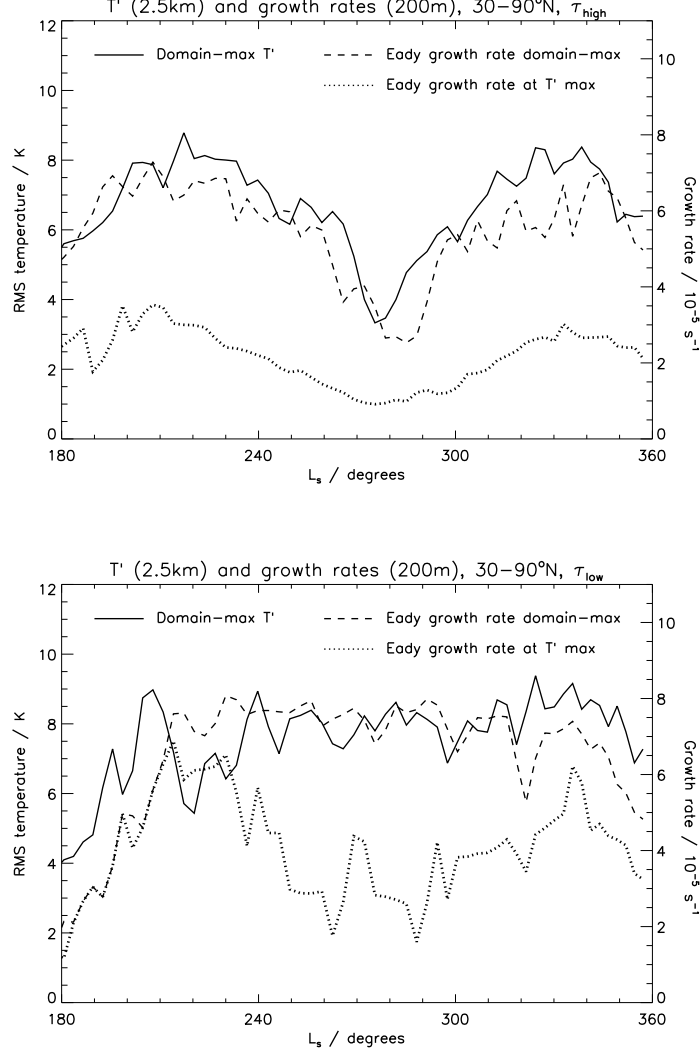


Fig. 9. As in Fig. 7, but for τ_{high} (top) and τ_{low} (bottom).

solstice value. By contrast, in τ_{low} , a $\sim 50\%$ solstitial reduction in local Eady growth rate does not lead to any reduction in peak eddy temperature. Meanwhile, the maximum growth rate in the domain reaches a seasonal maximum around solstice. The differences between τ_{low} and τ_{high} here support the idea that consistently high growth rates at higher latitudes can contribute to the maintenance of eddy activity in the midlatitudes.

To summarise the findings of the last two subsections, then, in each of τ_{MY24}^* ,

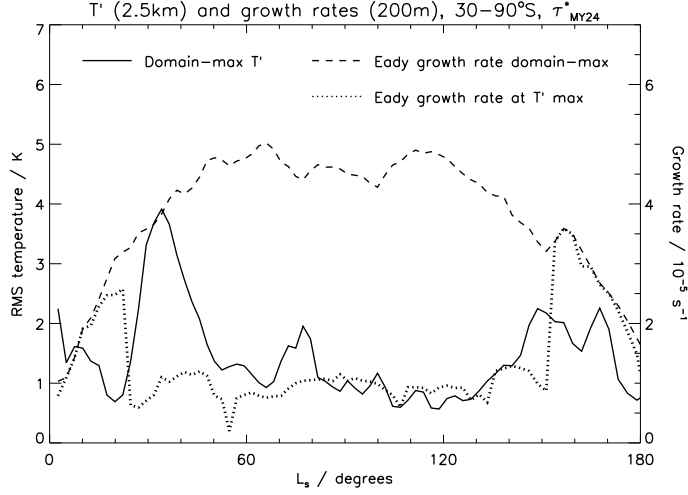


Fig. 10. As Fig. 7, but for 30–90°S during southern autumn and winter, in simulation τ_{MY24}^* .

τ_{MY24} and τ_{high} (as well as τ_{low}^* , not shown), the simulated solstitial pause forms due to a meridional tilting of the westerly jet: for the dust storm case this affects the whole jet and is due to the enhanced seasonal expansion of the PMOC, whereas for the more typical non-storm case, the equatorward shearing occurs primarily near the ground, and is greatly enhanced (in the model) by the presence of the polar hood clouds. In both cases, vertical wind shear and eddy growth rates at 60–80°N are affected, which leads to changes in eddy activity to the south of 60°N.

4.1.5 Southern hemisphere

The southern hemisphere winter jet, though weaker than in the north, also exhibits a poleward tilt over its full height, which is significant over $L_s \approx 30$ –150°. A strong reduction in windspeed poleward of 60°S near the ground is seen in the model at this time, and it is slightly enhanced when clouds are included. Clouds also reduce \bar{u} and $\frac{\partial \bar{u}}{\partial z}$ at the ground at 40–50°S, giving a

deeper solstitial pause in τ_{MY24}^* than in τ_{MY24} (Fig. 3). By contrast, the wind shear does not weaken around solstice in τ_{low} , and no pause is produced.

The southern hemisphere minimum in τ_{MY24}^* is seen despite consistently large Eady growth rates at $\sim 70^\circ\text{S}$ (Fig. 10, dashed line), in contrast to what was
385 seen in the northern hemisphere of τ_{low} (Fig. 9). Variation in RMS temperature is in rough agreement with the local growth rate (Fig. 10, dotted line), as at solstice both fall to around one third of their pre-solstice values, and both show post-solstice increases. The fact that, unlike in the northern hemisphere, the presence of larger Eady growth rates at higher latitudes does not sustain
390 eddy activity around solstice could perhaps be explained by the greater separation in latitude between the south polar cap edge and the peak growth rate region that exists around solstice (during $L_s = 60\text{--}140^\circ$, the peak temperature variance is located at $30\text{--}40^\circ\text{S}$ (c.f. $40\text{--}50^\circ\text{N}$), while the peak in growth rate remains at $60\text{--}70^\circ\text{S}$) or by the existence of less favourable conditions for the
395 equatorward propagation of disturbances. The baroclinic zone is much closer to the strong $60\text{--}70^\circ\text{S}$ growth rates during the pre- and post-solstice RMS temperature maxima, due to the smaller polar cap that exists away from solstice.

In the reanalysis of southern autumn and winter of MY25 (not shown), there
400 is a more straightforward link between growth rate and eddy activity, as the largest Eady growth rate remains close to the location of strongest eddy activity, and this growth rate agrees well with the resulting RMS temperature. The large solstitial values seen in Fig. 10 are not found in the reanalysis, perhaps because the peak moves equatorward rather than remaining at $60\text{--}70^\circ\text{S}$ (see
405 subsection 4.3).

4.2 Other eddy energy sources/sinks

In addition to the change to the basic state baroclinicity caused by the presence of water ice clouds, two other mechanisms were investigated, each of which could potentially affect the availability of transient eddy energy in the autumn and winter seasons, by introducing additional source or sink terms.

4.2.1 Barotropic shear

On Earth, Deng and Mak (2006) suggested that the Pacific midwinter suppression can be partly explained by barotropic eddy dissipation. On Mars, transient eddy energy is generated barotropically, and especially so around solstice (see P1, Fig. 6), which would appear to rule this out as a factor in the solstitial pause. However, there is a possible indirect route for barotropic shear to contribute to the suppression of near-surface eddies.

James and Gray (1986) noted that flow with a substantial barotropic component can become significantly stabilised with respect to baroclinic instability, due to deformation of any disturbances which should occur at a rate proportional to the horizontal wind shear, \bar{u}_y . By comparing this rate to the expected linear (Eady) growth rate, they suggested that instability is significantly reduced when

$$\frac{N\bar{u}_y}{f\bar{u}_z} \gtrsim 0.31 \quad (1)$$

James and Gray envisaged the barotropic governor as converting eddy kinetic energy (EKE) into zonal mean kinetic energy, and additionally reducing the rate of baroclinic conversion from eddy available potential energy (APE) to EKE. Model EKE generation rates for τ_{MY24} and τ_{MY24}^* (not shown) reveal

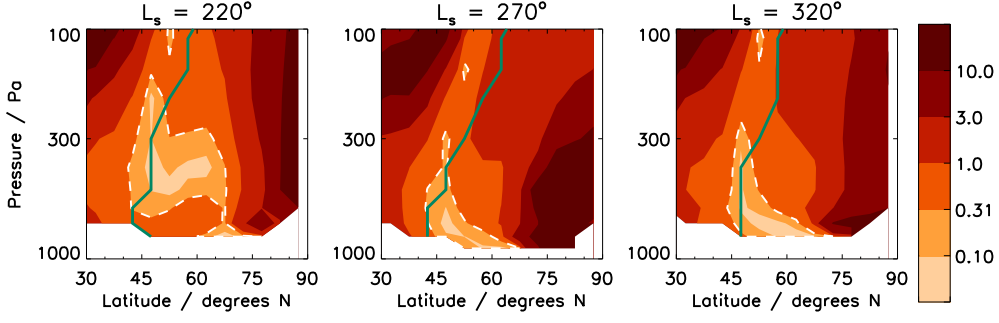


Fig. 11. The dimensionless quantity $\frac{N\bar{u}_y}{f\bar{u}_z}$, calculated for simulation τ_{MY24}^* , during autumn/winter in the northern hemisphere. The dashed white contour shows where this quantity is equal to 0.31, and the solid line marks the latitude of the strongest meridional temperature gradient at each model height level. The volume within the 0.31 contour should be ‘protected’ from the barotropic governor’s effect.

that barotropic conversion peaks around winter solstice in both cases, in the middle atmosphere and to a lesser degree close to the surface. Both the solstitial peak in barotropic generation and the minimum in baroclinic generation
 425 are amplified by the addition of clouds, suggesting a possible causal link between the two via the barotropic governor effect.

Fig. 11 shows that large portions of the lower atmosphere (those outside the dashed region) do satisfy condition (1), but at the latitude of strongest temperature gradient it is generally not satisfied near the surface, implying that
 430 the most unstable latitudes remain, to some extent, protected from any suppressive influence of the barotropic wind shear. Although the volume of the atmosphere in which the condition is not satisfied decreases around solstice, it remains largely intact at the surface. The seasonal variation of this volume, and of its position with respect to the peak temperature gradient, is very similar
 435 in the cloud and no-cloud (not shown) cases, which display rather different

solstitial pauses. Therefore, it would seem that the barotropic governor is not primarily responsible for the suppression near the surface during winter. An increase in the barotropicity of the atmosphere around solstice does seem to be important — it is seen consistently in each simulation that produced a
440 solstitial pause, caused either by the presence of clouds or by a dust-enhanced circulation around solstice — but the transition from baroclinic to barotropic conversion in the lower atmosphere can be seen in the reduction of vertical wind shear as presented in the previous subsection, and it does not seem necessary to invoke any additional suppression to explain the observed reduction
445 in eddy activity.

4.2.2 *Diabatic dissipation*

Chang (2001) found from a modelling study that a contribution to the mid-winter suppression on Earth could be provided by a diabatic dissipation of eddy APE, caused by surface sensible heat flux. Away from winter solstice,
450 he found that stronger generation of eddy APE from condensational heating was enough to overwhelm the dissipation by the surface heat flux, thereby increasing storminess.

Fig. 12 shows the rate of generation of eddy APE in the northern hemisphere of the two ‘MY24’ simulations, due to diabatic heating and through
455 baroclinic conversion from zonal-mean APE. The diabatic generation rate is (following Peixoto and Oort (1992)) $\int \Gamma \overline{Q'T'} dm$, where Q' is the eddy diabatic heating rate and Γ is an inverse static stability parameter defined as $\Gamma = -\left(\frac{\kappa\theta}{pT}\right)\left(\frac{\partial\tilde{\theta}}{\partial p}\right)^{-1}$, where $\tilde{\theta}$ is the domain-mean potential temperature on a particular pressure surface. The rate includes contributions from direct heat-

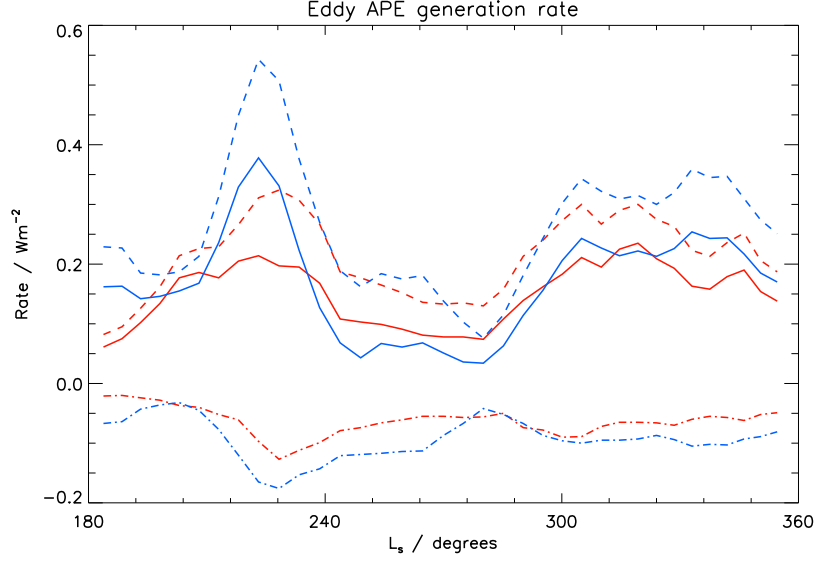


Fig. 12. Eddy available potential energy generation rates (in W m^{-2}), vertically integrated and averaged over $30\text{--}90^\circ\text{N}$, for τ_{MY24} (red) and τ_{MY24}^* (blue). Lines show contributions from baroclinic conversion (dashed) and diabatic heating (dash-dotted), and their sum (solid).

ing by dust, CO_2 and water ice clouds, and the surface sensible heat flux. The contribution from latent heat of condensation is several orders of magnitudes smaller, so is neglected. In both cases, diabatic effects (dash-dotted lines) dissipate eddy APE throughout the autumn/winter period, meaning that perturbations in radiative cooling are positively correlated with positive temperature perturbations, and thus act to damp baroclinic waves. Since most of the winter polar atmosphere receives little or no solar flux, this reflects the fact that warm regions tend to cool faster than colder regions.

Diabatic dissipation of eddy APE acts in opposition to, and primarily varies in response to, its generation baroclinically from zonal-mean APE (dashed lines), which peaks before and after solstice in both cases. The net result (generation – dissipation, solid lines) shows seasonal progressions that agree

with the respective near-surface temperature variances, with a lower net APE generation rate over $L_s = 240\text{--}300^\circ$ for τ_{MY24}^* . Over $\sim 240\text{--}270^\circ$, baroclinic conversion is at least as rapid for τ_{MY24}^* as for τ_{MY24} , and the lower net rate
475 is due to greater diabatic dissipation. However, for $\sim 270\text{--}300^\circ$, the lower net rate for τ_{MY24}^* is the result of reduced baroclinic conversion. From these domain-average figures, it appears that thermal damping caused by water ice clouds could be partly responsible for the deeper solsticial pause produced in τ_{MY24}^* , at least in late autumn ($L_s = 240\text{--}270^\circ$).

480 4.3 Topographic heterogeneity

With or without ice clouds, winter midlatitude and polar dynamics are undoubtedly influenced significantly by the substantial and complex topography that exists on Mars. Fig. 13 shows the difference between the zonal mean martian surface and the areoid (the martian equivalent of terrestrial sea level), and
485 the standard deviation of the surface height in the zonal direction. Most of the northern hemisphere slopes downward towards the pole, while the southern polar region slopes upward towards the pole. Blumsack and Gierasch (1972) suggested that the northern hemispheric surface, with the lower boundary sloping in the opposite sense to atmospheric isentropes, should reduce baroclinic instability and shift the most unstable mode to shorter wavelengths,
490 with the magnitude of the effect increasing with the steepness of the gradient. The meridional topographic slope is steeper over $35\text{--}50^\circ\text{N}$ than it is over $50\text{--}75^\circ\text{N}$, which would suggest that transient eddy activity in the northern hemisphere should be reduced if and when the baroclinic zone moves south of
495 $\sim 50^\circ\text{N}$. Using the temperature gradients of Fig. 3(d) as a guide, this progres-

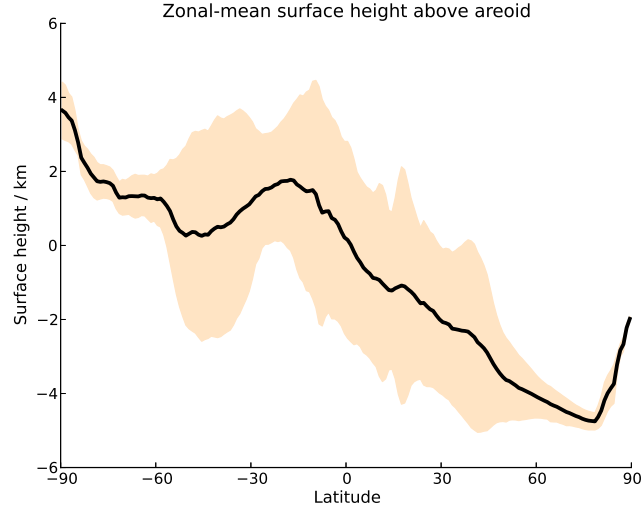


Fig. 13. Height of the zonally averaged martian surface above the areoid (km, solid line). Shading (mean $\pm 1\sigma$) indicates standard deviation in surface height around latitude circles.

sion to or beyond 50°N appears to coincide with the duration of the solsticial pause. A similar argument could be made for the southern hemisphere in the region north of 50°S , as the slope changes direction, although the meridional variation is more complicated than it is in the northern hemisphere.

500 The martian surface is also notably more zonally asymmetric in midlatitudes than it is poleward of 60° (note the varying spread in Fig. 13). Yu and Hartmann (1995) found that adding topographic features to a zonally symmetric surface results in a transfer of wave energy from transient to stationary waves, and from high to low frequencies. In their paper, ‘high frequency’ was de-
505 fined as those waves with a period of less than 7 days: therefore, it is possible that most martian transient waves, having relatively short periods, will be suppressed to some extent by the presence of the large zonal variations in topography that exist equatorward of 60° .

A critical latitude for suppression of instability can be hypothesised, situated
 510 at around $50\text{--}60^\circ$ in each hemisphere. Since the seasonal ice caps grow beyond
 this latitude around midwinter, a topographic contribution to the seasonal
 variation of eddy activity is possible, caused by one or both of the increase
 in meridional slope and the increase in zonal asymmetry equatorward of 60° .
 To investigate which, if either, of these effects exert an influence, two fur-
 515 ther model runs were performed. One simulation, $\text{zonav_}\tau_{MY24}^*$, used zonally
 averaged topography, while the other, $\text{zmrem_}\tau_{MY24}^*$, had the mean merid-
 ional slope of the topography removed by subtracting the zonal mean surface
 height from each gridpoint. Both were forced by the MY24 dust scenario and
 included radiatively active water ice clouds. The dust scenario was not zonally
 520 or meridionally smoothed; however, the scenario is defined relative to a refer-
 ence pressure level, so dust opacities do not retain a signal of the non-averaged
 martian surface, only of heterogeneities that existed in the opacity field at the
 reference level.

The results, which can be compared to those in Fig. 3(d), are shown in Fig.
 525 14. In $\text{zonav_}\tau_{MY24}^*$, transient temperature perturbations in the northern hemi-
 sphere are increased over those from τ_{MY24}^* throughout autumn and winter,
 and a solstitial pause is no longer produced; in fact, there is now a solstitial
 maximum. This is accompanied by a consistently large vertical wind shear (and
 therefore large baroclinic growth rate) at $60\text{--}75^\circ\text{N}$, and a secondary maximum
 530 at $10\text{--}30^\circ\text{N}$ (not shown). The region of peak temperature variance spreads
 further south, as far as $30\text{--}40^\circ\text{N}$, where previously there was a strong $s = 3$
 topographic forcing.

The difference in solstitial pause depth between τ_{MY24} and τ_{MY24}^* can, there-
 fore, be at least partially attributed to the influence of the clouds in shifting

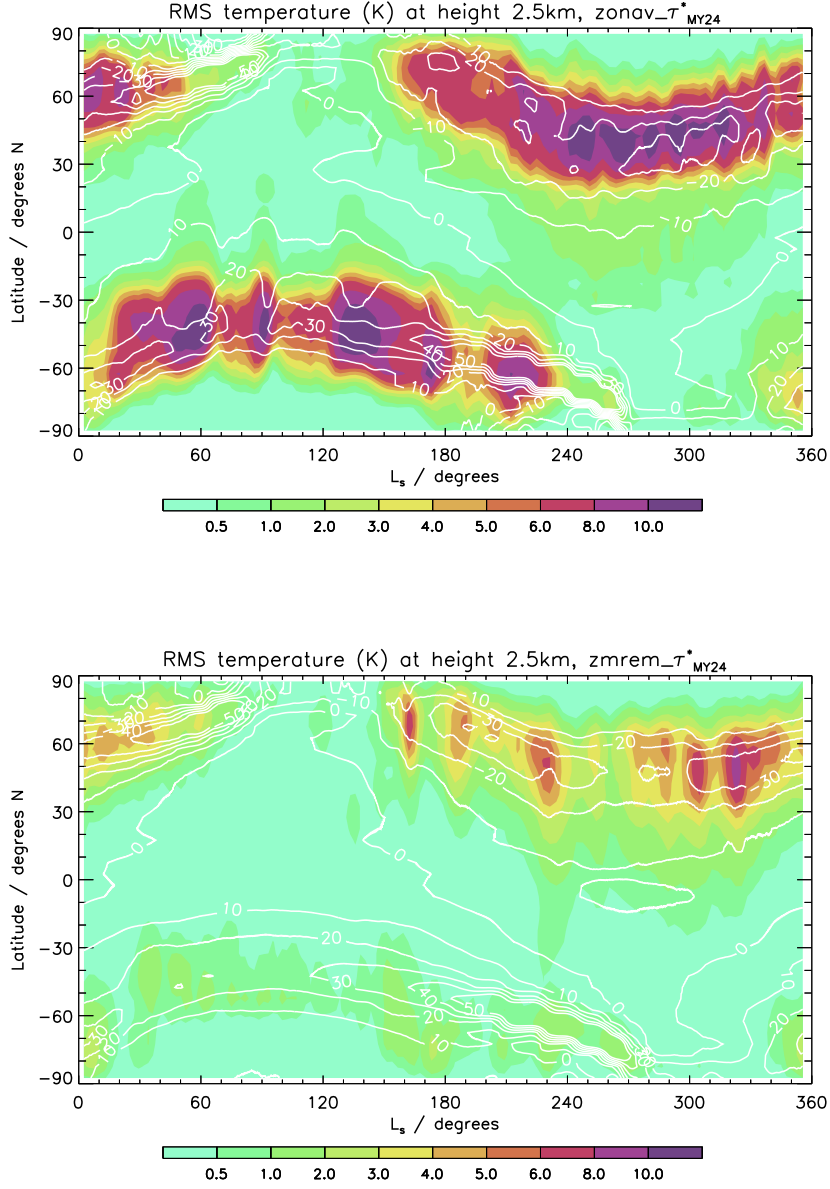


Fig. 14. Transient RMS temperature (K) at 2.5 km for runs using the MY24 dust scenario with radiatively active water ice clouds, and using zonally averaged topography ($\text{zonav-}\tau_{MY24}^*$, top), and topography with the mean meridional slope removed ($\text{zmrem-}\tau_{MY24}^*$, bottom). The meridional temperature gradient at 200 m (10^{-6} K m^{-1}) is overlaid in white contours.

535 the region of peak eddy activity equatorwards and further into the topograph-
ically asymmetric midlatitudes, where eddy growth seems to be impeded.

A dramatic change is seen in the southern hemisphere, where peak RMS tem-
peratures increase and are now of comparable magnitude to those in the north-
ern hemisphere. Several short maxima and minima are seen but transient ed-
540 dies remain fairly strong over $L_s = 30\text{--}150^\circ$ at $30\text{--}50^\circ\text{S}$, whereas in τ_{MY24}^*
they are largely absent from this region/season. It appears that, in reality, the
zonal topographic asymmetry at these latitudes (which is predominantly of
wavenumbers 1 and 3) is a major reason for the weakness of eddy activity in
the southern hemisphere, where RMS values $\gtrsim 4\text{ K}$ are rarely seen north of
545 60°S . In this simulation with zonally averaged topography, peak Eady growth
rates at 200 m at $60\text{--}70^\circ\text{S}$ (not shown) are much larger than for τ_{MY24}^* , though
in fact growth rates are increased throughout $40\text{--}80^\circ\text{S}$, and peaks in RMS
temperature appear to follow fluctuations in local growth rate at $\sim 45^\circ\text{S}$ for
much of the autumn/winter period.

550 The simulation $\text{zmrem-}\tau_{MY24}^*$ produces a marginally shallower minimum in
the northern hemisphere than τ_{MY24}^* , implying that an increased meridional
slope at $35\text{--}50^\circ\text{N}$ does (in τ_{MY24}^*) slightly suppress instability. More notably,
however, there is a general reduction in temperature variance in the northern
hemisphere: with the exception of the pre-solstice peak, Eady growth rates
555 are larger than they are in τ_{MY24}^* , yet weaker waves result throughout autumn
and winter. Increased growth rates are consistent with the removal of a lower
boundary that slopes downwards towards the pole, but these do not result
in stronger eddies, possibly a result of the decreased surface pressure in the
northern hemisphere, as the surface now sits higher relative to the areoid
560 than in τ_{MY24}^* , which implies a lower CO_2 frost point and greater temperature

inversions over the cold polar cap, making the atmosphere more statically stable. It would appear that this has a greater effect on transient eddies than does the change to the slope of the lower boundary.

The southern hemisphere sees very little change from τ_{MY24}^* , suggesting a
 565 greater degree of balance between the effects of changing surface pressure and changing meridional slope.

5 Discussion

5.1 *The role of topography*

From comparing the zonally averaged topography simulation $\text{zonav-}\tau_{MY24}^*$ to
 570 those described in earlier sections, it is clear that the zonal asymmetry of the midlatitude topography exerts a very strong influence on transient activity in both hemispheres. In the case of the southern hemisphere, the removal of this asymmetry increases model RMS temperature during autumn and winter by around a factor of two, bringing magnitudes close to those of the northern
 575 hemisphere, and implying that, although winter baroclinicity is slightly weaker in the southern hemisphere on account of Mars's orbital eccentricity, it is in fact the topographic zonal asymmetry that is primarily responsible for the relatively weak transient eddies observed in southern autumn and winter.

In the northern hemisphere, the absence of zonal variations in topography
 580 allows a large surface vertical wind shear to be maintained throughout autumn and winter at 60–70°N, where it would otherwise tend to reduce around solstice, as the polar ice cap edge extends into the more zonally heteroge-

neous midlatitude region. A secondary peak in wind shear and growth rate at 10–30°N is possibly also important in creating the solstitial maximum seen in zonally-averaged τ_{MY24}^* . These results suggest that while a decrease in atmospheric baroclinicity close to the surface is the direct explanation for the northern hemisphere solstitial pause, this arises more fundamentally due to the existence of zonally asymmetric topography in the midlatitudes.

Additional altered-topography runs could help to determine why the zonal asymmetry in topography suppresses waves before and after solstice so much more strongly in the southern hemisphere than in the northern hemisphere. The amplitudes of the zonal topographic variability are fairly similar in each of the two hemispheres, but perhaps crucially the zonal standard deviation is larger at around 50°S than at 50°N (Fig. 13). It may be necessary to look at the longitudinal patterns of energy generation, as in Kavulich Jr et al. (2013), rather than using zonal-mean diagnostics, to understand the specific impacts of topographic features in both hemispheres. This issue is pertinent to the occurrence of cap-edge dust storms (and potentially also larger storms) during the recent history of Mars, since the precession of the solar longitude of perihelion, which occurs with a period of ~ 50 kyr (Ward, 1974), is roughly equivalent to the flipping of the topography about the equator every ~ 25 kyr.

Further study of the zonal distribution of martian eddy activity could also reveal the similarities and differences that exist between the roles of topography in the solstitial pause and in the terrestrial midwinter suppression. In the latter case, interactions between the surface and low-level flow in one specific region, central Asia, have been suggested to reduce local baroclinic conversion so as to create a downstream solstitial minimum in eddy activity over the North Pacific in some years (Penny et al., 2010; Park et al., 2010). On Mars,

the reduction in eddy activity occurs to a similar degree at all longitudes (see
610 P1, Figs. 7 & 8), suggesting that an ‘upstream seeding’ model may not be appropriate. However, there may exist similarities between the martian situation and the seasonal variation in the eddy energy generation process that occurs over the central Asian mountains.

5.2 *Model dependence*

615 The effects of dust, ice cloud and topographic forcings have been diagnosed above in the UKMGCM, but it must be noted that other models show some variation in simulated solsticial pause depth. For example, Wang et al. (2013) show a pause of comparable magnitude to that of τ_{MY24}^* , from a simulation without radiatively active ice clouds and with a moderate increase in dust
620 opacity near solstice. Basu et al. (2006) also present a plausible solsticial pause, though it is unclear what dust loading was used for this simulation. It may be that the dynamics of these models differ from those of the UKMGCM, such that they simulate a sharper drop-off in near-surface vertical wind shear around solstice in a cloud-free run. Indeed, differences have been
625 seen between the two versions of the LMD/UK model (using either the LMD or UK dynamical core), under the same dust forcing.

Supporting the idea that the presence of water ice clouds contributes to the formation of the solsticial pause is the modelling work of Wilson (2011), who also simulated a deep solsticial pause only when radiatively active water ice
630 clouds were included. The addition of active clouds to the other models mentioned above can therefore be expected to increase the depth of the simulated pauses, which would bring them even closer to observations. However, the

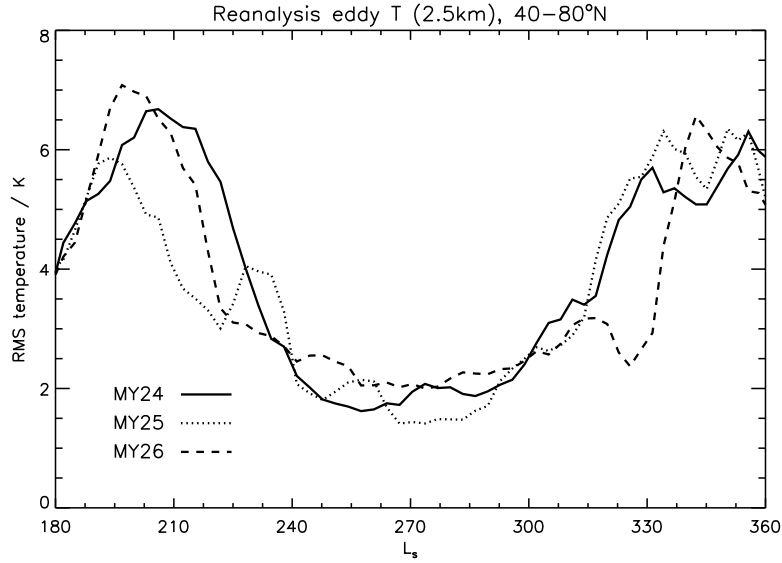


Fig. 15. Transient RMS temperature at 2.5 km, averaged between 40°N and 80°N, for MY24 (solid), MY25 (dotted) and MY26 (dashed), from the reanalysis.

magnitude of the impact of ice clouds may differ in other models, due to their stronger cloud-free simulated solsticial pauses.

635 5.3 Interannual variability in near-surface travelling waves

Our insight into the mechanisms behind the suppression of near-surface eddies in northern midlatitudes allows us to draw some conclusions regarding the interannual variability seen in northern hemispheric travelling waves in the TES reanalysis data, building on ideas discussed in Wang (2007). The three
640 years of transient eddy temperature data from the reanalysis can be seen in Figs. 1 and 2 of P1; here, Fig. 15 illustrates more clearly the variation in the length and depth of the northern hemisphere solsticial pause at the surface.

The earlier reduction in transient eddy activity seen in MY26 compared to MY24 followed the earlier development of the upper maximum and tilt of

645 the westerly jet due to the regional storm at $L_s = 210^\circ$ in MY26. Wang
 (2007) found that flushing storms, which disappeared upon the onset of the
 210° storm, did return briefly over $L_s = 230\text{--}250^\circ$ which, as can be seen
 from Fig. 1 of P1, was a period of moderate transient activity along the cap
 edge. She attributed the eventual cessation of flushing storms at $L_s = 250^\circ$ to
 650 the progression into the solstitial window, where the change in atmospheric
 state makes it impossible to sustain flushing storm activity. The particular
 aspect of this change that is relevant, according to the results shown here in
 previous sections, is the tilt of the westerly jet, particularly near the ground,
 which increases naturally around solstice, even without a substantial increase
 655 in global dust loading (Figs. 5, 6).

Conversely, a large dust storm may not fully suppress surface transient ed-
 dies if it occurs too early in the year. The pre-solstice peak in surface RMS
 temperature was lower in MY25 than in the other two years, but nevertheless
 values of ~ 6 K were present at $L_s = 200^\circ$, while the global dust storm was
 660 in progress. It seems that pre-210° is simply too early to force a sufficient tilt
 in the jet near the ground to fully suppress surface instability, even with a
 storm as large as the one in MY25. This fits with the results of simulation
 τ_{high} (Fig. 9), in which eddy activity was not reduced until after $L_s \approx 220^\circ$,
 despite an equally high dust loading during $L_s = 180\text{--}220^\circ$.

665 The post-solstice peak in eddy activity was delayed by $\sim 20^\circ$ in MY26 com-
 pared to the previous two years: this can be linked to the dramatic increase
 in dust loading seen towards the end of that year (the $L_s = 315^\circ$ dust storm),
 which led to a temporary reversal of the seasonal decrease in the magnitude
 of the polar warming, and associated tilt of the jet, that usually occurs at
 670 this time of year. When the late-winter storm occurred, the temperature at

(10 Pa, 60°N), which had been declining since $L_s = 270^\circ$, increased again, by ~ 10 K, as the PMOC expanded in response to the increase in dust loading. The tilt that this imparted to the jet allowed a renewed suppression of the near-surface vertical wind shear and therefore of the transient eddy activity, which had been increasing in line with the other two years until the occurrence of the storm.

Dust lifting in the northern midlatitudes therefore appears to possess a self-limiting quality, whereby the occurrence of a series of flushing storms is likely to inhibit the development of further such storms through the strengthening of the PMOC (Wang, 2007). Periods of flushing storm activity in early autumn and late winter seem certain to occur, however, because outside $L_s \approx 210\text{--}330^\circ$ it is very difficult to force the tilt in the lower part of the westerly jet that is needed to reduce near-surface growth rates. What is less certain is that there must always be a window where no flushing storms occur. Whether or not these conditions occur in a particular year may depend on there being a sufficiently large dust loading around northern winter solstice, resulting either from the occurrence of a sequence of flushing storms or more simply by an increase to the ‘background’ loading above some baseline level. The examination of further Mars years recently added to the reanalysis (Montabone et al., 2015) has the potential to shed further light on this matter.

Although no such increased dust loading can be invoked to explain the annual pause in the southern hemisphere, it is possible that the global circulation is altered in an equivalent way by the presence of the thick tropical cloud belt around solstice (Madeleine et al., 2012), affecting the tilt of the southern winter jet. This would explain the presence of a southern hemisphere solstitial pause in τ_{low}^* (Fig. 3(b)) but not in τ_{low} (Fig. 3(a)).

6 Conclusions

The martian solsticial pause, as presented in P1, has been convincingly simulated by the UKMGCM, when run with a realistic dust loading and the
700 inclusion of radiatively active water ice clouds. Both increased dustiness and ice clouds alter the thermal structure of the polar atmosphere, decreasing the zonal wind and its vertical shear near the surface at mid- and high latitudes around winter solstice, causing a decrease in baroclinic growth rates.

Baroclinic eddy activity near the surface in both hemispheres is reasonably
705 well predicted by linear instability theory. In all model runs that produce a solsticial pause, and to a lesser extent in the reanalysis of MY24, variation in transient eddy amplitudes is accompanied by variation in Eady growth rate, driven by changes in near-surface vertical wind shear as well as in the background state (f , N). Quantitative agreement was found between variations in
710 growth rate and the amplitude of temperature perturbations, though in general it is necessary to consider the growth rate at 60–70° latitude rather than at the location of the eddy maximum, in midlatitudes. There is evidence that a high wind shear at these high latitudes can sustain eddy activity at midlatitudes, though the effect was seen not to apply in the southern hemisphere,
715 perhaps due to different meridional flow conditions near the surface.

The barotropic governor effect was found not to play a significant role in the martian solsticial pause in either hemisphere. Diabatic dissipation, in the form of radiative damping, may play a modest role, in concert with modulations in the local Eady growth rate itself. Zonal variations in midlatitude topography
720 exert a strong control on near-surface transient eddies, reducing their ampli-

tude dramatically in the southern hemisphere throughout autumn and winter, and around winter solstice in the northern hemisphere. It appears, therefore, that the zonal topographic asymmetry in midlatitudes is the ultimate cause of the solstitial pause in both hemispheres.

725 **Acknowledgements**

The authors thank Don Banfield and an anonymous reviewer for their careful and helpful reviews. We acknowledge useful discussions with John Wilson and Luca Montabone. This work was funded by the UK Science and Technology Facilities Council.

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